

Solving radiative transfer with line overlaps using Gauss Seidel algorithms.

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Accepted 2008 June 13. Received 2008 April 30

ABSTRACT

Context. The improvement in observational facilities requires refining the modelling of the geometrical structures of astrophysical objects. Nevertheless, for complex problems such as line overlap in molecules showing hyperfine structure, a detailed analysis still requires a large amount of computing time and thus, misinterpretation cannot be dismissed due to an undersampling of the whole space of parameters.

Aims. We extend the discussion of the implementation of the Gauss–Seidel algorithm in spherical geometry and include the case of hyperfine line overlap.

Methods. We first review the basics of the short characteristics method that is used to solve the radiative transfer equations. Details are given on the determination of the Lambda operator in spherical geometry. The Gauss–Seidel algorithm is then described and, by analogy to the plan–parallel case, we see how to introduce it in spherical geometry. Doing so requires some approximations in order to keep the algorithm competitive. Finally, line overlap effects are included.

Results. The convergence speed of the algorithm is compared to the usual Jacobi iterative schemes. The gain in the number of iterations is typically factors of 2 and 4 for the two implementations made of the Gauss–Seidel algorithm. This is obtained despite the introduction of approximations in the algorithm. A comparison of results obtained with and without line overlaps for N_2H^+ , HCN, and HNC shows that the $J = 3 - 2$ line intensities are significantly underestimated in models where line overlap is neglected.

1. Introduction

Our understanding of the evolution of the molecular complexity in astrophysical media relies both on the quality of the observations and on their interpretation. The increasing spatial resolution that will be provided by the new generation of ground– and space–based telescopes (Herschel, ALMA, JWST) will enable finer geometrical structures to be inferred in the objects under study. Hence, the best way to interpret the origin of the molecular line emission and to derive physical and chemical conditions is to use non local radiative transfer calculations with realistic geometries adapted to the angular resolution. While the Sobolev approximation can be safely used when a large velocity gradient is present, interstellar clouds often show linewidths that are thermal in nature or that involve low–velocity fields, so the use of such an approximation is rather doubtful, and a widely used method in the field of molecular astrophysics is the Monte–Carlo technique (Bernes, 1979). This method is particularly attractive due to its conceptual simplicity and because it is easily implemented to treat multidimensional geometries. Nevertheless, an inherent drawback of the method is that it suffers from numerical noise and presents a slow convergence rate when optically thick lines are modelled. Although line overlaps can be easily implemented (Gonzalez-Alfonso & Cernicharo, 1993), computing time becomes very expensive when treating a case such as N_2H^+ where the two nitrogen atoms contribute to the hyperfine structure.

Since the 90s, various techniques have been introduced to solve the coupled set of radiative transfer and statistical equilibrium equations. These methods lead to substantial gains in term of calculation speed and accuracy of the solutions. Among them, the preconditioning of statistical equilibrium equations entails drastic improvement in the convergence properties of the iterative scheme for problems with large optical depths (see Rybicki & Hummer, 1991, 1992). A crucial issue is that the Lambda iteration scheme does not always converge to the true solution for large optical depths (see Mihalas, 1978; Auer, 1991). Another progress consisted in the introduction by Trujillo Bueno & Fabiani Bendicho (1995) of the Gauss–Seidel algorithm for the solution of radiative transfer problems.

The description of this method was first made for the case of one-dimensional plane-parallel stellar atmospheres (Trujillo Bueno & Fabiani Bendicho, 1995; Paletou & Léger, 2007) and was further generalised to the case of two and three-dimensional radiative transfer problems with multilevel atoms in Cartesian coordinates (Fabiani Bendicho et al., 1997; Léger et al., 2007) and to the case of scattering line polarization (Trujillo Bueno & Manso Sainz, 1999). It was then described (Asensio Ramos & Trujillo Bueno, 2006) and used (Cernicharo et al., 2006) in spherical geometry. Although the geometry of interstellar clouds cannot always be assumed to be as simple as spherical, it enables a first guess more appropriate than plan–parallel geometry in order to model centrally condensed clouds. In this article, we give more details on the implementation in spherical geometry, and we extend the discussion to the

case of line overlap. We find that in spherical geometry, the requirements of the original Gauss–Seidel algorithm cannot be fulfilled without increasing the computational time (even in the case of isolated lines). To address this problem, we introduced several approximations into the algorithm. Interestingly, despite these approximations, the usual increase in reaching convergence in terms of number of iterations persists. Moreover, these approximations do not entail false convergence, in comparison to results obtained within the Λ iteration scheme.

Several molecules used to trace the physical conditions in molecular clouds harbour a hyperfine structure that complicates the analysis of the observed emission (see e.g. Daniel et al., 2007). Indeed, the influence of the overlap is usually not considered when interpreting observations despite the fact that for a molecule with hyperfine structure, the line blending increases with the rotational quantum number. For high rotational transitions, all the hyperfine lines are usually overlapped in astrophysical objects. Previous works have shown that, for HCN, line overlaps play an important role in the pumping of molecular levels and have to be considered in order to interpret the relative intensities of hyperfine transitions (cf. Guilloteau & Baudry, 1981; Gonzalez-Alfonso & Cernicharo, 1993). One goal of this paper is to extend the discussion to other molecules, i.e. HNC and N_2H^+ , and to derive the general radiative transfer effects that are common to the different species.

This article is organised as follows. In section 2.1, we first overview the short characteristics (SC) method. In section 2.2, the main equations used in the preconditioning of statistical equations are given. In section 2.3, the implementation of the Gauss–Seidel algorithm in spherical geometry is discussed. In section 3, a test case for H_2CO gives the main properties of the various algorithms. Finally, in section 4, the influence of introducing line overlap in the calculations is discussed for the N_2H^+ , HCN, and HNC molecules.

2. Theory

2.1. Short characteristics in spherical geometry

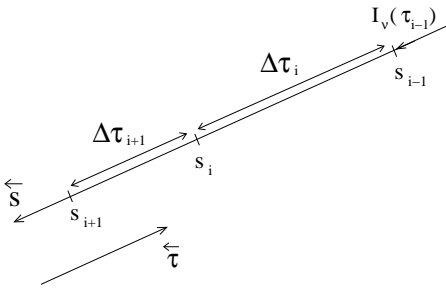


Fig. 1. Transfer along a short characteristic.

To solve the transfer equation, we use the SC method introduced by Olson & Kunasz (1987). Here, we briefly overview the method and give the main equations in order to introduce quantities that will be used in the following discussion.

If one considers the radiation propagation along a ray that is discretized (cf. Figure 1), the specific intensity between two consecutive points varies as (Olson & Kunasz, 1987)

$$I_v(\tau_i) = I_v(\tau_{i-1}) e^{-\Delta\tau_i} + \int_{\tau_{i-1}}^{\tau_i} S_v(\tau) e^{\tau_i - \tau} d\tau$$

$$\Leftrightarrow I_i = I_{i-1} e^{-\Delta\tau_i} + \Delta I_i \quad (1)$$

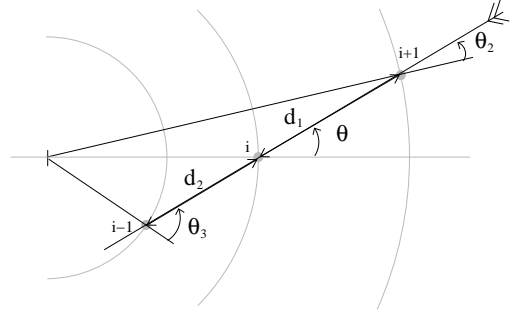


Fig. 2. Geometric quantities in spherical geometry.

Assuming a parabolic trend for the source functions with respect to the opacity of the line, we can derive the intensity analytically as follows:

$$\Delta I_i = \alpha_i S_v(\tau_{i-1}) + \beta_i S_v(\tau_i) + \gamma_i S_v(\tau_{i+1}) \quad (2)$$

where the expressions of the coefficients α_i , β_i , and γ_i are given by Olson & Kunasz (1987) and depend on $\Delta\tau_i$ and $\Delta\tau_{i+1}$. These latter quantities are derived assuming that the absorption coefficients $\kappa_v(s)$ vary linearly between two consecutive grid points.

In spherical geometry, the radiation field is fully characterised introducing two variables, r and θ . Numerically, one has to adopt a discretized space for these two variables. At a given radius r_i and for an angle θ_i , the specific intensity is known within the SC scheme from the geometrical quantities represented in Fig. 2. In the case $\theta_i < \pi/2$, we find

$$\begin{aligned} d_{1,i} &= -r_i \cos \theta + \sqrt{r_{i+1}^2 - r_i^2 \sin^2 \theta_i} \\ \theta_{2,i} &= \arccos \left(\frac{d_{1,i}^2 + r_{i+1}^2 - r_i^2}{2 d_{1,i} r_{i+1}} \right) \\ d_{2,i} &= r_i \cos \theta_i - \sqrt{r_k^2 - r_i^2 \sin^2 \theta_i} \\ \theta_{3,i} &= \arccos \left(\frac{r_i^2 - r_k^2 - d_{2,i}^2}{2 d_{2,i} r_k} \right) \end{aligned} \quad (3)$$

$$\text{with } k \equiv \begin{cases} i+1 & \text{if } \theta_i = \pi/2 \\ i & \text{if } \theta_i \in]\pi/2; \xi_{i,i-1}[\\ i-1 & \text{if } \theta_i \geq \xi_{i,i-1} \end{cases} \quad (4)$$

$$(5)$$

The angle $\xi_{m,n}$ is defined according to

$$\xi_{m,n} = \pi - \arccos \sqrt{1 - \left(\frac{r_n}{r_m} \right)^2} \quad (6)$$

and stands for the angle sustained by the layer of radius r_n seen from the layer of radius $r_m > r_n$.

The quantities $d_{1,i}$ and $d_{2,i}$ are used to derive $\Delta\tau_i$ and $\Delta\tau_{i+1}$. The angle $\theta_{2,i}$ is used to interpolate the incoming intensity. Moreover, the angle $\theta_{3,i}$ is used in the determination of $\Delta\tau_{i+1}$ when a radial velocity field is introduced. For the case $\theta_i > \pi/2$, the above formulae also apply, using the angle $\pi - \theta_i$ and then inverting respectively $d_{2,i}$ with $\theta_{3,i}$ and $d_{1,i}$ with $d_{2,i}$.

2.2. Accelerated Lambda Iteration

Following Olson et al. (1986), rather than the lambda iteration scheme (LI), we use an alternative scheme known as the accelerated (or approximated) lambda iteration (ALI) since this one presents better convergence properties. In that scheme, the specific intensity obeys

$$I_{\nu\Omega} = (\Lambda_{\nu\Omega} - \Lambda_{\nu\Omega}^*)[S_{\nu}^+] + \Lambda_{\nu\Omega}^*[S_{\nu}] + \mathcal{J}_{\nu\Omega}^+ \quad (7)$$

where $\Lambda_{\nu\Omega}$ stands for the exact Lambda operator and $\Lambda_{\nu\Omega}^*$ for an approximated Lambda operator (ALO). In this expression, the indexes correspond to the frequency ν and to the direction of propagation $\Omega \equiv (r, \theta)$. The quantities indexed by $+$ are determined using the level populations obtained at the previous iteration. The value $\mathcal{J}_{\nu\Omega}^+$ corresponds to the external radiation field which is transmitted to the point under consideration. From Eq. 7 the LI scheme would be obtained letting $\Lambda_{\nu\Omega}^* \equiv 0$.

Several studies aiming at determining the most effective ALO to be used so that the convergence is reached in the fastest way have been conducted. Non diagonal Lambda operators were found to be efficient in reducing the number of iterations but at a computational cost per iteration higher than for a diagonal Lambda operator. Finally, the use of a diagonal ALO seems to lead to the fastest convergence rates (see e.g. Puls & Herrero, 1988; Olson et al., 1986).

When we use the diagonal part of the full lambda operator for $\Lambda_{\nu\Omega}^*$, we find that instead of solving the usual statistical equilibrium equations (SEE), one has to solve the system (Rybicki & Hummer, 1991)

$$\begin{aligned} 0 = & \sum_{j>i} [n_j A_{ji}(1 - \bar{\Lambda}_{ji}^*) - (n_i B_{ij} - n_j B_{ji}) \bar{\mathcal{J}}_{ji}^{\text{eff}}] \\ & - \sum_{j<i} [n_i A_{ij}(1 - \bar{\Lambda}_{ij}^*) - (n_j B_{ji} - n_i B_{ij}) \bar{\mathcal{J}}_{ij}^{\text{eff}}] \\ & + \sum_{j \neq i} [n_j C_{ji} - n_i C_{ij}] \end{aligned} \quad (8)$$

with

$$\bar{\mathcal{J}}_{ij}^{\text{eff}} = \bar{\mathcal{J}}_{ij}^+ - \bar{\Lambda}_{ij}^* \times S_{ij}^+ \quad (9)$$

In the above equation, S_{ij} is the usual frequency-independent source function for an isolated molecular line. The term $\bar{\mathcal{J}}_{ij}^+$ corresponds to the radiation field averaged over frequency and angle. The above expressions apply with or without any source of background continuum radiation. When there is a source of background continuum radiation, the source function is given by (Rybicki & Hummer, 1991)

$$S_{\nu} = r_{ij} S_{ij} + (1 - r_{ij}) S_c \quad (10)$$

with

$$r_{ij} = \frac{\kappa_{ij}(\nu)}{\kappa_{ij}(\nu) + \kappa_c(\nu)} \quad (11)$$

where $\kappa_{ij}(\nu)$ is the absorption coefficient due to the molecular transition and $\kappa_c(\nu)$ is the continuum absorption coefficient. The averaged Lambda operator is given by

$$\bar{\Lambda}_{ij}^* = \frac{1}{4\pi} \int_{\Omega} \int_{\nu} \phi_{ij}(\nu) \Lambda_{\nu\Omega}^* r_{ij}^+ d\nu d\Omega \quad (12)$$

where $\phi_{ij}(\nu)$ is the intrinsic line shape function.

2.2.1. Overlap of hyperfine lines

The overlap between hyperfine lines is treated as described by Rybicki & Hummer (1992). Here we give a brief overview of the main equations. In the case of overlapping lines, the source function without background continuum radiation is given by

$$S_{\nu} = \frac{1}{\kappa(\nu)} \sum_{ij} \kappa_{ij}(\nu) S_{ij} \quad (13)$$

where S_{ij} is the usual frequency independent source function encountered in the case of an isolated line and κ_{ν} is the total absorption coefficient at frequency ν given by

$$\kappa(\nu) = \sum_{ij} \kappa_{ij}(\nu) \quad (14)$$

Within the SC scheme, these quantities are used to perform the analytical integration along a characteristic using Eqs. 1 and 2. As explained in Rybicki & Hummer (1992), the preconditioning of the statistical equilibrium equations can be made in different ways. In the present work, we use the preconditioning with the same transition. This corresponds to eq. 2.23 of Rybicki & Hummer (1992). Rewriting this latter equation in terms of the Lambda operator rather than the Ψ operator which acts on emissivities and considering that a diagonal ALO acts like a simple multiplication leads to the same form of the preconditioned statistical equilibrium equations as given by Eq. 8. The only difference is that the averaged diagonal ALO is now given by

$$\bar{\Lambda}_{ij}^* = \frac{1}{4\pi} \int_{\Omega} \int_{\nu} \phi_{ij}(\nu) \frac{\kappa_{ij}(\nu)}{\kappa(\nu)} \Lambda_{\nu\Omega}^* d\nu d\Omega \quad (15)$$

We note that the full preconditioning strategy might be more adapted to the case of the hyperfine structure where lines are close in frequency. Thus, the present choice is motivated by the SEE equations remaining unchanged and are still given by Eq. 8.

In the appendix, the way the Lambda operator is constructed is presented and the differences with the plan-parallel case are emphasised.

2.3. Gauss-Seidel

The Gauss-Seidel (GS) iterative scheme was introduced in radiative transfer by Trujillo Bueno & Fabiani Bendicho (1995). The superiority of this algorithm was then shown, by comparison to Jacobi-based ones, in order to rapidly reach convergence even for the complex case of multilevel atoms in two and three-dimensional Cartesian geometry (see Fabiani Bendicho et al., 1997). Recently, a detailed description of the implementation of the GS algorithm for multilevel atomic (or molecular) case was presented by Paletou & Léger (2007). So far, the GS algorithm has been extensively described in plan-parallel geometry. The description of the implementation in spherical geometry has been done Asensio Ramos & Trujillo Bueno (2006), and here we extend the discussion made in this article. Before, we will review the main attributes required by the algorithm.

At each iteration, the grid is first swept from the outermost layer until the innermost one (say, from $k = N \rightarrow 1$), and during this process, part of the averaged radiation field is computed, i.e. for $\theta \in [0; \pi/2]$. Then, once the innermost

layer is reached, the grid is swept in the other direction. This leads to the computation of the averaged radiation field that corresponds to $\theta \in [\pi/2; \pi]$. Within the Jacobi iterative scheme, all the grid is swept until reaching the outermost layer, and then statistical equilibrium equations are solved. Within the GS iterative scheme, once the averaged radiation field is known for a layer, the statistical equilibrium equations are solved for it. At this stage, various updates have to be made to account for the newly derived populations. Using parabolic interpolation, this implies correcting both the outward and inward radiation field respectively, for the current and the next grid point. These corrections are made so that when deriving the populations for a layer k , the populations considered are those obtained during the current iteration for the layers $i < k$ and those corresponding to the previous iteration for the layers $i \geq k$, noted $\{n_j^{new}\}_{i < k}$ and $\{n_j^{old}\}_{i \geq k}$.

We now describe how the GS algorithm should be implemented in spherical geometry and how we have implemented it in practice. In what follows, we note I_k^{in} , the intensity $I_k(\theta)$ for $\theta \in [0; \pi/2]$ and I_k^{out} the intensity $I_k(\theta)$ for $\theta \in [\pi/2; \pi]$.

2.3.1. Upward inversion

Let us consider first that the grid has been swept from the outermost layer to the innermost one. This leads to the computation of $\{I_k^{in}\}_{k \in [1; N]}$. Once the innermost layer is reached, we can compute the intensities I_1^{out} since the incoming intensities are given by I_1^{in} (If there is a central source of continuum radiation, the intensities are an input parameter.). We then know \bar{J} , and we can invert the statistical equilibrium equations for this layer. Before going to the point $k = 2$, there are three corrections that in principle have to be done to conform with the requirements of the original GS algorithm:

- an update of I_1^{in} which is computed with n_1^{new} and n_2^{old}
- an update of I_1^{out}
- an update of I_2^{in} computed with n_1^{new} and n_2^{old} and n_3^{old} .

Note that the first correction should be performed because the intensities I_1^{in} are used to derive the incoming intensities when determining I_1^{out} . However, this correction is not necessary in practice since it influences the convergence rate just slightly (cf. section 3). Once these corrections are done, we go forward in the propagation.

Suppose we reach the point k where new populations are obtained. Let us now consider the corrections that are needed to have an exact GS iterative scheme. Since new populations have been obtained for the layer k , a requirement of the algorithm is that, at the point $k + 1$, the populations should be determined with $\{n_j^{new}\}$ for $j \in [1; k]$ and $\{n_j^{old}\}$ for $j \in [k + 1; N]$. Now, consider the way I_k^{out} depends on the various populations and so, through the dependance on the incoming intensities. For this purpose, we refer to Fig. 3, which shows the dependencies of intensities. We see that to obtain the intensities I_k^{out} , it is necessary to know the inward intensities I_k^{in} , as well as the outward intensities I_{k-1}^{out} , which in turn depend on I_{k-2}^{out} , etc. Finally, we see that fully accounting for the new population n_k^{new} in the computation of I_k^{out} would require first to perform the propagation along the inward direction, leading to the updated set $\{I_j^{in}\}_{j \in [1; k+1]}$. Then, the propagation would be done again along the outward direction so that the set $\{I_j^{out}\}_{j \in [1; k]}$ would effectively be given considering the new populations n_k^{new} . In practical applications, such a process would

lead to computational times longer than those obtained within the Jacobi iterative scheme. Moreover, it should be noted that this problem does not arise in plan-parallel geometry: in that case, once the population is updated at point k , the computation can be performed at point $k - 1$ and will take exactly the new populations into account (cf. Fig. 3), since the radiation field at the innermost boundary is an input parameter of the model.

To make the algorithm computationally competitive, we proceed as follows. By analogy with the plan-parallel case, a way to implement the Gauss-Seidel algorithm would be not to account for the influence of the new populations n_k^{new} on the incoming intensities I_{k-1}^{in} and I_k^{in} when we compute the outgoing intensities I_k^{out} . This would lead to performing

- an update of I_{k-1}^{out} using the new populations at points $k - 2, k - 1$, and k using I_{k-2}^{out} and I_{k-1}^{in} for the incoming intensities, calculated with the old populations.
- an update of I_k^{out} using the populations n_{k-1}^{new} , n_k^{new} , and n_{k+1}^{old} . As for the previous correction, the incoming intensities I_k^{in} are given by the old populations.
- an update of I_{k+1}^{in} using the populations n_k^{new} , n_{k+1}^{old} , and n_{k+2}^{old} .

As previously said, the first correction is in fact not necessary for speeding up the convergence rate. Actually, the mixing of *old* and *new* populations for the point k reduces slightly the convergence rate (cf. section 3). Finally, we only perform the two latest corrections, which is equivalent to the corrections described by Asensio Ramos & Trujillo Bueno (2006).

2.3.2. Upward and downward inversion

As explained above, within the usual GS scheme and for each iteration, the grid is first swept from $k = N$ to $k = 1$ and then from $k = 1$ to $k = N$, and during this second step, statistical equilibrium equations are solved. As mentioned by Trujillo Bueno & Fabiani Bendicho (1995) and Trujillo Bueno & Manso Sainz (1999), a way to increase the convergence speed is to invert the statistical equilibrium equations during both the upward and downward sweeping of the grid.

During the downward propagation and when the iterative process has been initiated, we already know the intensities $\{I_j^{out}\}_{j \in [1; N]}$. Thus, when reaching point k , we directly compute the averaged radiation field for $\theta \in [0; \pi]$. First we compute I_k^{in} and then I_k^{out} using I_k^{in} and I_{k-1}^{out} as incoming intensities. Here we note that the intensities I_{k-1}^{out} in fact involve a mixing of the populations obtained during the two last iterations. This is because the first correction described in the previous section is not performed and that the intensities are saved after the second correction. Finally, since the intensities I_k^{out} are computed at this stage, there is only one correction needed so that, after inversion, we perform an update of I_k^{in} using the populations n_{k+1}^{new} , n_k^{new} , and n_{k-1}^{old} . During the downward propagation, we have to save the intensities I_k^{in} calculated with the new populations since these intensities will be used during the upward propagation. The upward sweeping of the grid is similar to the one described in the previous section.

3. Numerical results

The numerical implementation was checked against the results obtained with the code described by Gonzalez-Alfonso & Cernicharo (1993). Various models

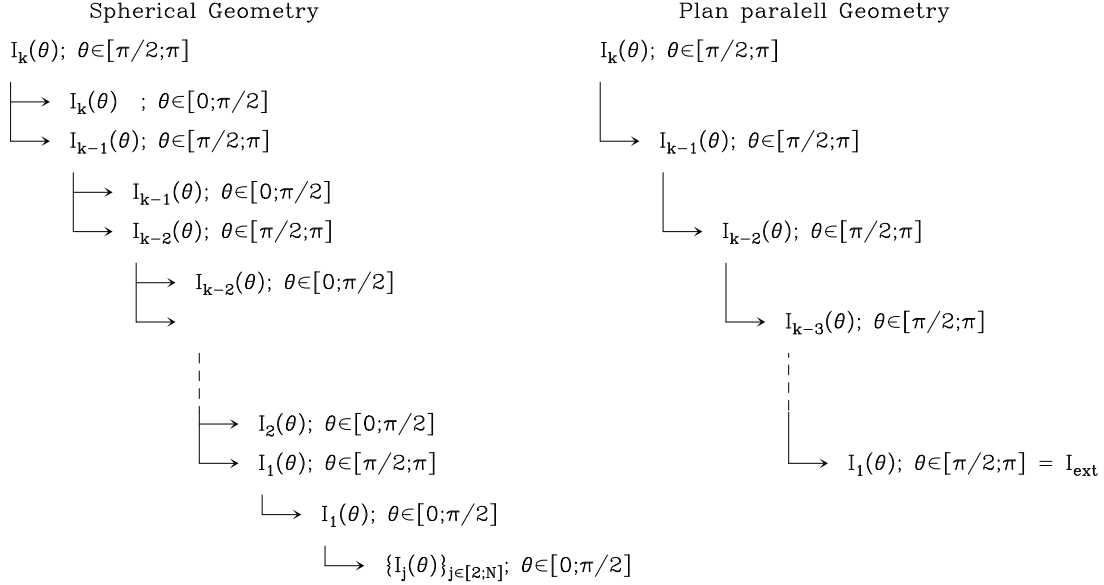


Fig. 3. Dependence of the intensities.

were used in the comparison, covering different physical conditions, as well as low, moderate, and large optical depths. The differences in the derived level populations ranged from a tenth of a percent to a few percent when integrated optical depths were higher than ~ 100 . This corresponds to the expected accuracy of numerical radiative transfer codes as explained in van Zadelhoff et al. (2002).

To accelerate the convergence we used the ALI method described in Sect. 2.2. In spherical geometry, the diagonal of the full Lambda operator is given by Eqs. B.10, B.11, and B.12. However, determining the exact diagonal is time consuming: first, it requires the evaluation of the opacities $\Delta\tau_{m,m+l}^+$ that depend on all the internal shells crossed by the rays and second, it requires interpolating in angle the coefficients c_{m+l}^+ to obtain them at the values θ_{m+l}^+ . In practice, we thus chose not to keep the corresponding terms and we approximated the diagonal of the Lambda Operator just keeping, for $\theta > \pi/2$,

– if $m \geq K + 1$:

$$\Lambda_{m,m} = \beta_m(\theta_m^-) + \gamma_{m-1}(\theta_{m-1}^-) e^{-\Delta\tau_{m,m-1}^-} \quad (16)$$

– if $m = K$:

$$\Lambda_{m,m} = \alpha_m(\theta_m^-) + \beta_m(\theta_m^-) + [\beta_m(\theta_m^+) + \gamma_m(\theta_m^+)] e^{-\Delta\tau_{m,m}^+} \quad (17)$$

For $\theta \leq \pi/2$, we conserve the terms given by expressions B.1 and B.2.

These terms permit us to accelerate the rate of convergence notably by comparison to the LI scheme. Moreover, the approximation made in deriving the diagonal of the ALO does not induce false convergence. Both aspects can be seen in Fig. 4 where the maximum difference with respect to the converged solution is plotted. The model considered here corresponds to a sphere of uniform physical parameters with $n(\text{H}_2) = 10^4 \text{ cm}^{-3}$ and temperature $T = 25 \text{ K}$. The sampling consists in ~ 500 spatial grid points and 20 angles. This has been found to insure a convergence for the level populations better than 0.1 % for the first 12 levels. The molecule considered is ortho- H_2CO , and 14 energy

levels were included in the calculation. The molecular parameters were taken from the references given in Table 1. Since no analytical solution exists for this problem, we used as reference populations those obtained with the LI scheme. The error relative to the converged solution for the considered grid, noted C_e , is defined according to Eq. 20 of Auer et al. (1994).

Figure 4 also shows the rate of convergence of the different algorithms in terms of the iteration number. For the GS algorithms, two curves are shown: the grey one corresponds to the case where the incoming intensities are corrected (i.e. correction 1 of Sect. 2.3.1) and the black one corresponds to the case where this correction is omitted. We see that, for this model, the ALI scheme speeds up the convergence by a factor 1.3 in comparison to the LI scheme. The GS down and GS up-down enabled us to reduce the number of iterations by factors of 2.0 and 4.4 respectively by comparison to the ALI scheme. In Sect. 2.3, the various corrections to be done for these two algorithms were discussed. For each correction, the code calls a formal solver that takes the source functions of the lines as input and outputs the averaged radiation field and Lambda operator, as well as the intensities. Thus, considering the number of calls to the formal solver and in the different schemes, we have for each layer 2 calls in the ALI scheme, 4 calls in the GS down scheme, and 6 calls in the GS up-down scheme. The number of calls gives the relative time per iteration of each algorithm. In practice, the GS schemes implementation can be improved by performing some bookkeeping, in particular for the incoming intensities. These lead to times per iteration, in the GS down and GS up-down schemes, which are respectively 1.45 and 2.3 times greater than in the ALI scheme. These times were corrected for the time spent in the code in internal storage on the hard drive. Such a procedure is done in our implementation of the GS schemes while not for the ALI one and makes the GS scheme $\sim 20\%$ slower than permitted by the algorithmic. Moreover, these longer times are compensated for by the reduced number of iterations required (see above). Thus, for the current implementation, the GS down and GS up-down respectively reduce the total calculation time by factors of 1.4 and 1.9. The numerical implementation of the GS scheme could

molecule	rate coefficients	dipole moment	energy levels
H ₂ CO	Green (1991)	Fabricant et al. (1977)	Müller et al. (2005)
HCN	Monteiro & Stutzki (1986)	Ebenstein & Muentner (1984)	Müller et al. (2005)
HNC	Monteiro & Stutzki (1986)	Blackman et al. (1976)	Müller et al. (2005)
N ₂ H ⁺	Daniel et al. (2005)	Havenith et al. (1990)	Caselli et al. (1995)

Table 1. References for the molecular parameters used in this work.

still be improved to gain some computing time. The corrections in spherical geometry are formally similar to the one in plan-parallel geometry, and for this case, Paletou & Léger (2007) find that the time per iteration of the GS down scheme was only 1.3 longer than that of the ALI scheme. Our code has been written in a very modular and structured way. By decreasing the number of internal calls to routines and functions, we could still gain computing time and approach the 1.3 value of plan-parallel geometry.

Nevertheless, the GS schemes become more advantageous if, rather than parabolic interpolation, we use linear interpolation. In that case, there are fewer corrections to carry out so that the time per iteration of the GS down and GS up-down scheme are only 1.4 and 1.7 that of the ALI scheme. The drawback is that if using linear interpolation, one needs to use more grid points in order to converge to the true solution. Nevertheless, for the model considered here, we find that the differences in level populations are inferior to 1%. In Fig. 5, the line parameters of the most opaque lines are plotted for the two types of interpolations and the results are indistinguishable on the scale of the graph.

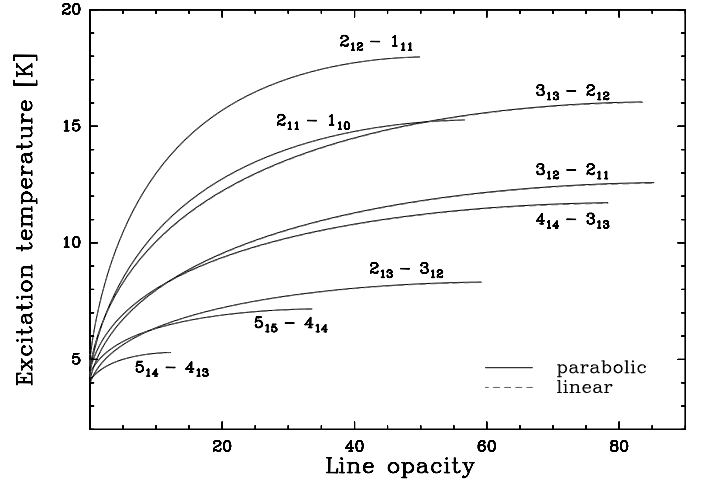


Fig. 5. Excitation temperature versus opacity for a model considering ortho-H₂CO.

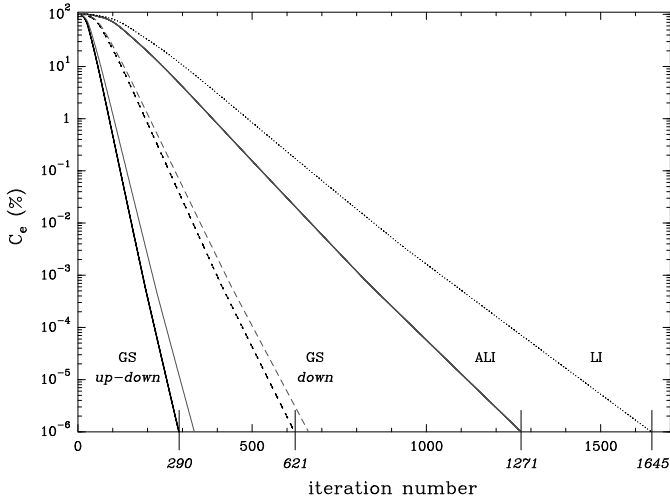


Fig. 4. Convergence properties of the various algorithms in terms of number of iterations.

We note that the various algorithms discussed here can still be improved if combined with the Ng acceleration technique (Ng, 1974). For the various models tested, we found that it typically leads to a reduction by a factor of 3 in the number of iterations. Alternatively, the successive over-relaxation technique can be used Trujillo Bueno & Fabiani Bendicho (1995).

4. Overlap of hyperfine lines

The numerical implementation of the overlap was checked according to the results reported by Gonzalez-Alfonso & Cernicharo (1993). For the molecules

considered in what follows the molecular parameters are summarised in Table 1. Figure 6 shows the variation in excitation temperature versus line opacity for the three HCN hyperfine lines associated with the $J = 1 - 0$ rotational transition. The physical parameters of the models were taken from Gonzalez-Alfonso & Cernicharo (1993). The line parameters are given with (solid lines) and without (dashed lines) introduction of the overlap. By comparison to the results reported in Figs. 1a–1d of Gonzalez-Alfonso & Cernicharo (1993), we see that the qualitative behaviour of T_{ex} is correctly reproduced for all the hyperfine lines when the overlap is introduced. Nevertheless, a quantitative comparison of the line parameters, with and without overlap, shows differences in the order of 5–10%. These differences are due to the choice of grid parameters, especially the number of radial grid points, adopted to reach convergence.

Figure 7 shows the influence of the overlap on the modelization of the hyperfine lines associated with the 3 lowest N₂H⁺ rotational transitions. This molecule is particularly interesting for inferring the physical conditions of the densest parts of cold dark clouds. Indeed, it was found that this molecule does not suffer from strong depletion for H₂ volume densities below $\sim 10^6$ cm⁻³. The model parameters are identical to those reported in Daniel et al. (2007) for the source L63. In this figure, we see that the $J = 1 - 0$ line remains mainly unchanged with differences in the line fluxes lower than 5%. The differences increase when considering higher rotational lines. For the $J = 3 - 2$ line, including the overlap increases the line flux by ~ 20 –30%. This can be understood as follows. For a group of blended hyperfine lines, the average radiation field seen by each line is greater when considering the overlap, since more molecules can interact with the radiation. Consequently, the pumping of the upper level of each transition is more efficient, which results in an enhancement

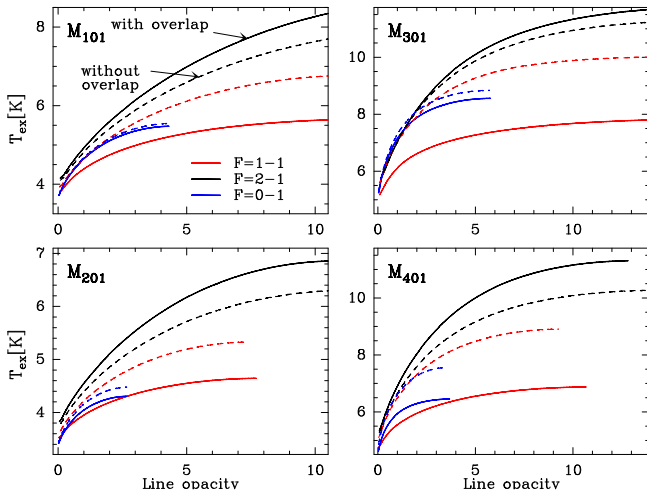


Fig. 6. Line parameters obtained for the models described in Gonzalez-Alfonso & Cernicharo (1993)

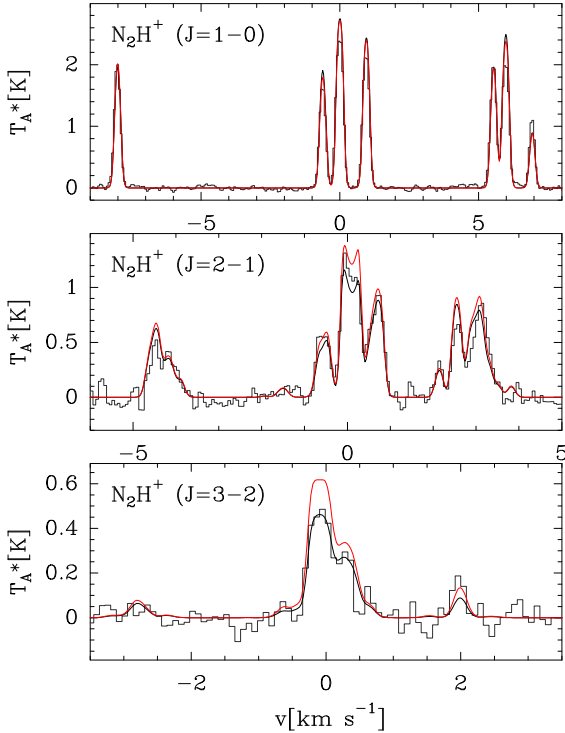


Fig. 7. Effect of the overlap on the line profile of L63. The black line correspond to a model without overlap and the red lines correspond to models that include overlap.

of the populations of the highest rotational levels. Therefore the overlap changes the physical parameters reported in Daniel et al. (2007). Mainly, H_2 volume densities and/or temperatures were overestimated in this study.

The effect induced by line overlap obviously depends on individual hyperfine line opacities, as well as on the frequency offset between the lines. Figure 8 shows the emerging intensities for the three lowest rotational lines of HCN, HNC, and N_2H^+ . The models consist of a uniform density core of diameter $1''$. The density is $n(H_2) = 10^5 \text{ cm}^{-3}$, the temperature $T=10 \text{ K}$, and the turbulence is set to $v_{turb} = 0.15 \text{ km.s}^{-1}$. For each molecule, calculations are performed at abundances of 10^{-10} , $2.5 \cdot 10^{-9}$, and $2.5 \cdot 10^{-8}$. First, it can be seen that, at the lowest abundance,

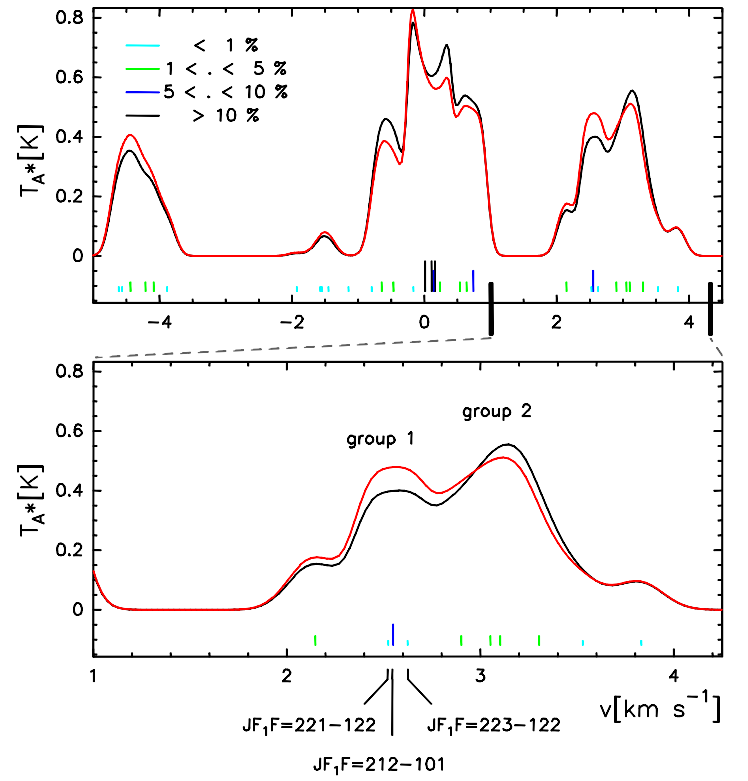


Fig. 9. Effect of overlap on relative line intensities. The black curve corresponds to a model that includes line overlap and the red line to a model without overlap, the latter being scaled by the integrated intensity ratio.

the effect introduced by the overlap is weak for N_2H^+ and HCN. The effect is stronger for HNC due to closer hyperfine frequencies. Second, we find that the effect introduced by the overlap is approximately proportional to the column density, as long as individual line opacities ≤ 1 . In the present calculations, this is found for the $J = 3 - 2$ lines of each species for the 2 lowest abundances. For these lines, we see that the integrated intensity ratios of the calculations with and without overlap are approximately proportional to the abundance ratio. In a more optically thick regime, the increase in the integrated intensity ratio depends less on the column density and finally reaches a saturation regime where the intensity ratio starts to be affected by self-absorption effects. This can be seen in Fig. 8 for the $J = 2 - 1$ and $J = 1 - 0$ lines. Finally, the error introduced in the interpretation of observations when neglecting the line overlap can be especially important for HNC. Indeed, we see in Fig. 8 that, for conditions typical of dark clouds, including line overlap leads to an intensity increase of up to a factor 2. Moreover, for similar abundances, the overlap differentially affects HCN and HNC and thus, neglecting it may lead to overestimating the $X(HNC)/X(HCN)$ ratio.

Another effect introduced by the overlap is to modify the relative intensities of the hyperfine components of a given rotational transition. Figure 9 shows the $J = 2 - 1$ line of N_2H^+ for a model with abundance $2.5 \cdot 10^{-9}$. The line obtained without overlap is represented multiplied by a factor of 1.54, which corresponds to the integrated intensity ratio obtained previously. In this figure the position in frequency of the different hyperfine components is also represented as are their line strengths. These are given in percent by reference to the value obtained by summation over all the hyperfine components. Considering the red

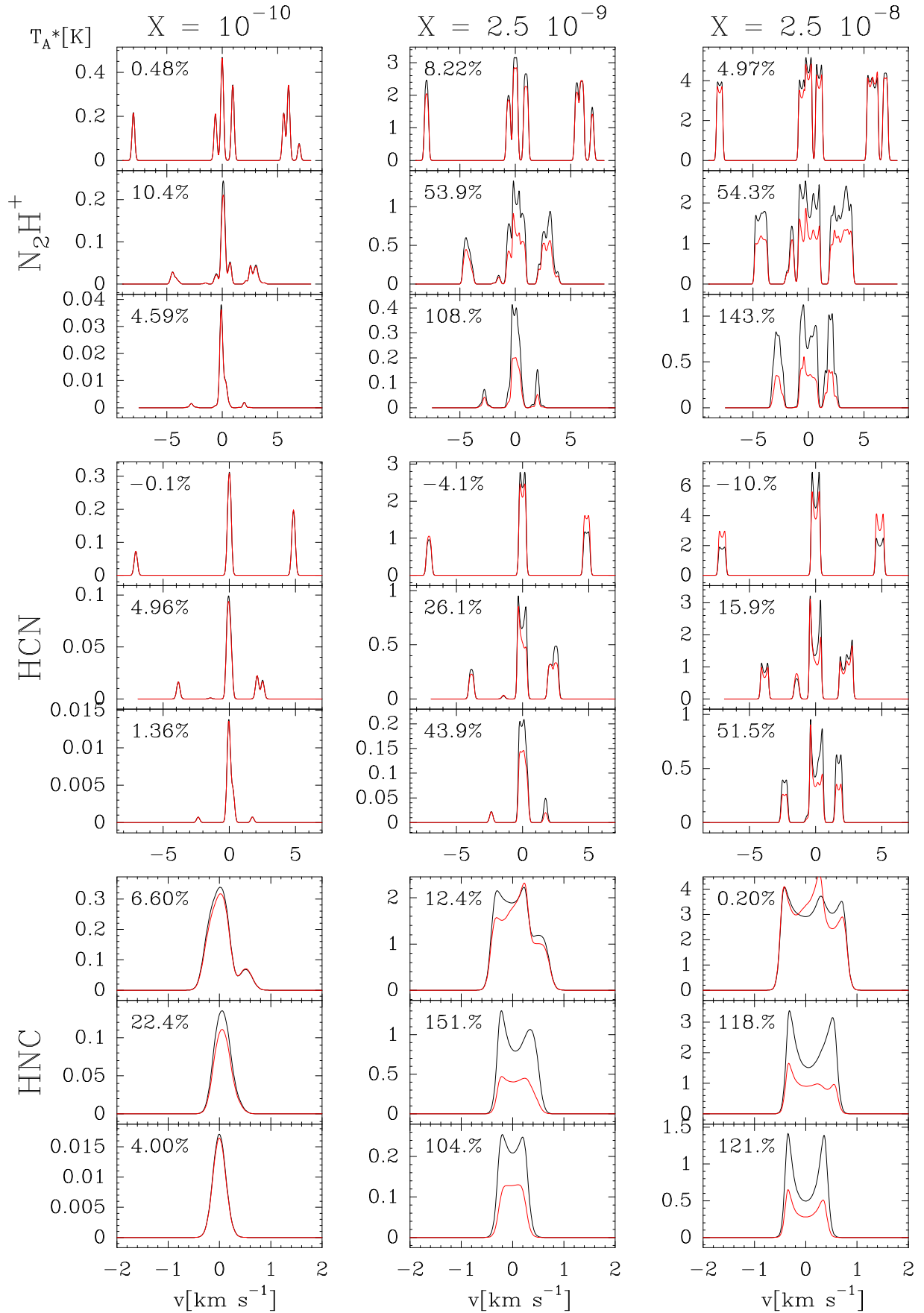


Fig. 8. Effect of overlap on the N_2H^+ , HCN, and HNC $J = 1 - 0$, $2 - 1$, and $3 - 2$ lines. Red lines correspond to models without introduction of the overlap and black lines to models considering line overlap. The differences between integrated intensities are given for each line. The intensity scale is antenna temperature and are obtained using the standard values of HPBW and beam efficiency of the 30m IRAM telescope.

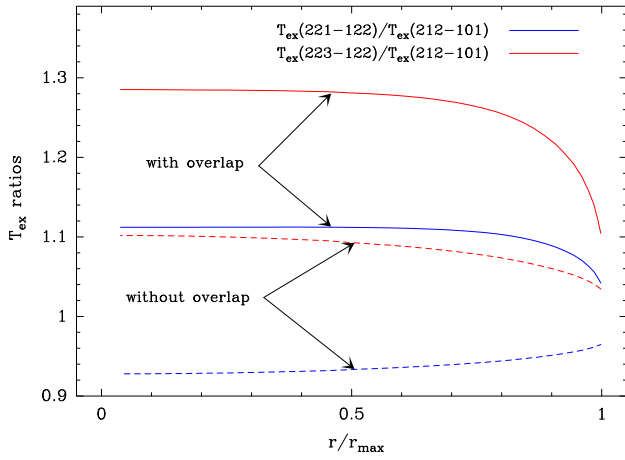


Fig. 10. Excitation temperatures ratios for overlapping lines with different line strengths.

satellite, we see that, when including the overlap, the intensity associated with the strongest transition (group 1 at $v \sim 2.5$ km s^{-1}) is reduced by comparison to the group of weaker lines situated at $v \sim 3$ km s^{-1} (group 2). This effect on the profile for the N_2H^+ $J = 2 - 1$ red satellite agrees with the observations reported in Daniel et al. (2007). Two effects yield to the current emerging profile. The main effect is introduced by the number of strong transitions that overlap in each group. In group 1, the line strengths of the two weakest transitions are low enough so that there is mostly one line that contributes to the emission. On the other hand, group 2 corresponds to 4 lines with equivalent line strengths. Thus, for this group of lines, the pumping of the upper levels is enhanced producing an increase in the emergent intensities. Another effect of the overlap is to create a flow of population that depends on the relative line strengths of the overlapping lines. Figure 10 shows excitation temperature ratios between the transitions of group 1 obtained with (solid lines) and without (dashed lines) introduction of the overlap. We see in this figure that the overlap enhances the excitation temperatures of the 2 weakest lines. This means that the levels $JF_1F = 223$ and $JF_1F = 221$ are more efficiently populated when considering the overlap due to the presence of the $JF_1F = 212 - 101$ transition that is close in frequency. This can be understood when considering the way the average radiation field is obtained when solving the radiative transfer. For the strongest line, the presence of weak hyperfine components close in frequency does not have a strong influence since the opacities/intensities in the wings of the intrinsic line profiles are low. The average radiation field is thus only slightly modified. On the other hand, for a weak line that is close to the frequency of a strong line, the average radiation field is increased. This will pump the corresponding energy levels and result in an increase in intensity for the weakest lines.

5. Conclusions

We have presented technical details concerning the determination of the Lambda operator. Moreover, we gave details on the implementation of the Gauss–Seidel algorithm in spherical geometry. Finally, the numerical code was applied to the treatment of line overlap. The main conclusions are:

1. The diagonal of the Lambda operator, as obtained within the short characteristics method, has to be approximated in practical applications. The use of approximated diagonal terms still enable speeding the convergence. Moreover, the tests

carried out did not show false convergence by comparison with LI calculations.

2. The implementation of the Gauss–Seidel algorithm in spherical geometry also requires some approximations. These are introduced by analogy with the plan–parallel case.
3. The convergence rate of the GS down and GS up–down are similar to what is obtained in plan–parallel geometry, i.e. they reduce the number of iterations to reach convergence by factors of 2 and 4 respectively. The gain in terms of computing time strongly depends on the numerical implementation, and in the present code, the GS up–down reduces the calculation time by a factor ~ 2 by comparison to the ALI scheme.
4. Line overlaps can lead to a large increase in the predicted intensities, the effect being greater for higher rotational lines. Especially, differential effects for HCN and HNC can alter the estimate of the abundance ratio $X(HNC)/X(HCN)$, and models where line overlap is neglected will overestimate it.

Acknowledgements. Part of this work was supported by the AYA2006-14876, ESP2004-00665, ESP2007-65812-C02-01, ESO n 010103050002, Molecular Universe MRTN-CT-2004-51230, ASTROCAM S-0505 ESP-0237 projects. The authors thank F. Paletou and L. Léger for the help provided concerning the Gauss–Seidel algorithm. We also thank J.R. Pardo for his useful revision of the present manuscript.

References

- Asensio Ramos, A., & Trujillo Bueno, J. 2006, EAS Publications Series, 18, 25
Auer, L. 1991, NATO ASIC Proc. 341: Stellar Atmospheres - Beyond Classical Models, 9
Bernes, C. 1979, A&A, 73, 67
Blackman, G.L., Brown, R.D., Godfrey, P.D. and Gunn, H.I. 1976, Nature 261, 395.
Caselli, P., Myers, P. C., & Thaddeus, P. 1995, ApJ, 455, L77
Cernicharo, J., Goicoechea, J. R., Pardo, J. R., & Asensio-Ramos, A. 2006, ApJ, 642, 940
Daniel, F., Dubernet, M.-L., Meuwly, M., Cernicharo, J., & Pagani, L. 2005, MNRAS, 363, 1083
Daniel, F., Cernicharo, J., Roueff, E., Gerin, M., & Dubernet, M. L. 2007, ApJ, 667, 980
Ebenstein, W. L., & Muentner, J. S. 1984, J. Chem. Phys., 80, 3989
Fabiani Bendicho, P., Trujillo Bueno, J., & Auer, L. 1997, A&A, 324, 161
Fabricant, B., Krieger, D., and Muentner, J.S. 1977, J. Chem. Phys. 67, 1576.
Gonzalez-Alfonso, E., & Cernicharo, J. 1993, A&A, 279, 506
Green, S. 1991, ApJS, 76, 979
Guilloteau, S., & Baudry, A. 1981, A&A, 97, 213
Havenith, M., Zwart, E., Leo Meerts, W., & Ter Meulen, J. J. 1990, J. Chem. Phys., 93, 8446
Léger, L., Chevallier, L., & Paletou, F. 2007, A&A, 470, 1
Mihalas, D. 1978, San Francisco, W. H. Freeman and Co., 1978. 650 p.,
Monteiro, T. S., & Stutzki, J. 1986, MNRAS, 221, 33P
Müller, H. S. P., Schlöder, F., Stutzki, J., & Winnewisser, G. 2005, Journal of Molecular Structure, 742, 215
Ng K.C., 1974, J. Chem. Phys., 61, 2680
Olson, G. L., Auer, L. H., & Buchler, J. R. 1986, Journal of Quantitative Spectroscopy and Radiative Transfer, 35, 431
Olson G.L. and Kunasz P.B. 1987, JQSRT, 38, 325
Rybicki, G. B., & Hummer, D. G. 1991, A&A, 245, 171
Rybicki, G. B., & Hummer, D. G. 1992, A&A, 262, 209
Paletou, F., & Léger, L. 2007, JQSRT, 103, 57
Trujillo Bueno, J., & Fabiani Bendicho, P. 1995, ApJ, 455, 646
Trujillo Bueno, J., & Manso Sainz, R. 1999, ApJ, 516, 436
Trujillo Bueno, J. 2003, in Stellar Atmosphere Modelling, eds. I. Hubeny, D. Mihalas, K. Werner, ASP Conf. Series Vol. 288, 551
Hauschildt, P.H. 1994, JQSRT, 47, 433
Puls, J., & Herrero, A. 1988, A&A, 204, 219
Auer, L., Bendicho, P. F., & Trujillo Bueno, J. 1994, A&A, 292, 599
van Zadelhoff, G.-J., et al. 2002, A&A, 395, 373

Appendix A: Discretized Lambda operator: plan-parallel geometry

As pointed out by Rybicki & Hummer (1991) the Lambda operator is in fact an approximation of the true Lambda operator since it depends on the discretization of the structure of the object. Thus, the Lambda operator is a square matrix whose dimension corresponds to the number of spatial grid points.

Within the SC scheme, the Lambda operator can be constructed easily. As a first step, from Eqs. 1 and 2 one finds that along a characteristic

$$I_i = I_0 e^{-\tau_i} + \underbrace{\sum_{k=0}^{i-1} \Delta I_{i-k}}_{\tilde{\Lambda}^i} e^{-\Delta\tau_{i,k}} \quad (\text{A.1})$$

where $\Delta\tau_{i,j} = \tau_i - \tau_j$. To obtain a compact expression, we adopt from now on the notation $a_i^{-1} = \alpha_i$, $a_i^0 = \beta_i$, and $a_i^1 = \gamma_i$ for the terms that appear in Eq. 2. Introducing it in the previous expression, we find

$$\tilde{\Lambda}^i = \sum_{k=0}^{i-1} \left(\sum_{l=-1}^1 \tilde{b}_{k+l}^{-l} e^{-\Delta\tau_{i,k+l}} \right) S_v(\tau_k) \quad (\text{A.2})$$

$$\text{with } \tilde{b}_{k+l}^{-l} = \begin{cases} a_{k+l}^{-l} & \text{if } 0 < k+l \leq i \\ 0 & \text{otherwise} \end{cases} \quad (\text{A.3})$$

In the above expression, the source functions are given at optical depths that correspond to the intersection of the ray with the discretized structure of the cloud. In order to construct the Lambda operator, it is then necessary to introduce a notation where the source functions are given as a function of grid points.

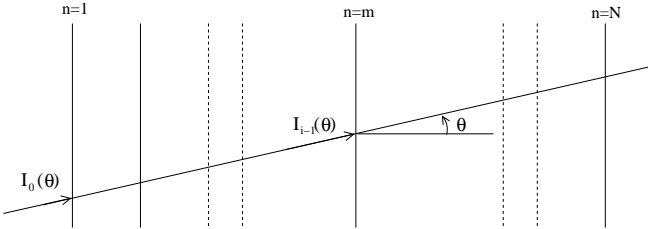


Fig. A.1. Notation adopted in plan-parallel geometry.

In plan-parallel geometry, the values for $S_v(\tau_k)$ are readily assigned to values of the source functions for each layer. From Fig. A.1, we have in the case $\theta \in [0; \pi/2]$ that the point $k=0$ in expression A.2 corresponds to the layer $n=1$. The point $k=i-1$ corresponds to the layer $n=m$. In the above expression, we thus first introduce the variable $n = k+1$ and then a notation where all the quantities are indexed according to the layer, i.e.

$$\begin{aligned} S_v(\tau_{n-1}) &\rightarrow S_v^n \\ \tilde{b}_{n-1+l}^{-l} &\rightarrow b_{n+l}^{-l} \\ \Delta\tau_{n-1,n-1+l} &\rightarrow \Delta\tau_{n,n+l}^- \end{aligned} \quad (\text{A.4})$$

The condition A.3 transforms as

$$b_{n+l}^{-l} = \begin{cases} a_{n+l}^{-l} & \text{if } 1 < n+l \leq m+1 \\ 0 & \text{otherwise} \end{cases} \quad (\text{A.5})$$

where the coefficient a_{n+l}^{-l} keeps the same definition given above, but is now labelled according to the layer index. It leads to

$$\tilde{\Lambda}^{i-1} = \Lambda^m = \sum_{n=1}^{m+1} \underbrace{\left(\sum_{l=-1}^1 b_{n+l}^{-l} e^{-\Delta\tau_{m,n+l}^-} \right)}_{\Lambda_{m,n}} S_v^n, \quad (\text{A.6})$$

In this expression, $\Lambda_{m,n}$ stands for the element of the m^{th} line and n^{th} column of the Lambda operator when $\theta \in [0; \pi/2]$.

Equivalent transformations can be done in the case $\theta \in [\pi/2; \pi]$. Defining N so that $n = N - k$, identifying $i = N - m$ and introducing the transformation

$$\begin{aligned} S_v(\tau_{N-n}) &\rightarrow S_v^n \\ \tilde{b}_{N-(n-l)}^{-l} &\rightarrow b_{n-l}^{-l} \\ \Delta\tau_{N-n,N-(n-l)} &\rightarrow \Delta\tau_{n,n-l}^- \end{aligned} \quad (\text{A.7})$$

with

$$b_{n+l}^{-l} = \begin{cases} a_{n+l}^{+l} & \text{if } m \leq n+l < N \\ 0 & \text{otherwise} \end{cases} \quad (\text{A.8})$$

leads to

$$\Lambda^m = \sum_{n=m-1}^N \left(\sum_{l=-1}^1 b_{n+l}^{+l} e^{-\Delta\tau_{m,n+l}^-} \right) S_v^n. \quad (\text{A.9})$$

This final expression defines the element $\Lambda_{m,n}$ when $\theta \in [\pi/2; \pi]$.

Appendix B: Discretized Lambda Operator: spherical geometry

In spherical geometry, the same procedure of the plan-parallel case can be followed to obtain the components of the Lambda operator. In this case, however, all the shells are not necessarily crossed for a given direction, and furthermore, some shells are crossed two times before reaching the current point.

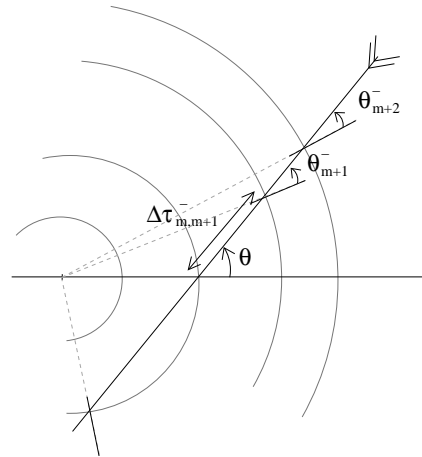


Fig. B.1. Definition of geometrical quantities.

We consider a cloud discretized in N concentric shells. In the case $\theta \in [0; \pi/2]$, the transformations are similar to the second case examined in the previous section with two exceptions. First, in the SC frame, the coefficients b_{n+l}^{+l} depend on an angle defined in a coordinate system associated to the layer $n+l$. This dependence was not outlined in the plan-parallel case since all

the coefficients along a characteristic are at a unique angle value, namely θ . In what follows, we introduce the notation θ_{n+l}^\pm , which refers to the angle to be taken into account in the evaluation of b_{n+l}^\pm (see Fig. B.1). The superscript in this notation distinguishes acute (-) and obtuse (+) angles.

Secondly, depending on the value of θ , the source function corresponding to the point $i+1$ can be associated either to shell m or to shell $m-1$. In this case, we obtain for $\theta \in [0; \pi - \xi_{m,m-1}]$

$$\Lambda^m = \sum_{n=m-1}^N \left(\sum_{l=-1}^1 b_{n+l}^\pm(\theta_{n+l}^\pm) e^{-\Delta\tau_{m,n+l}^\pm} \right) S_v^n, \quad (\text{B.1})$$

and, for $\theta \in [\pi - \xi_{m,m-1}; \pi/2]$

$$\Lambda^m = \sum_{n=m}^N \left(\sum_{l=-1}^1 b_{n+l}^\pm(\theta_{n+l}^\pm) e^{-\Delta\tau_{m,n+l}^\pm} \right) S_v^n + \gamma_m(\theta_m^-) S_v^m. \quad (\text{B.2})$$

In the above expressions, we have introduced the notation $\Delta\tau_{m,n}^\pm$. A ray that intersects the shells m and n defines two segments. The opacity along the longest one defines $\Delta\tau_{m,n}^+$, and the opacity along the smallest one defines $\Delta\tau_{m,n}^-$ (see Figs. B.1 and B.2).

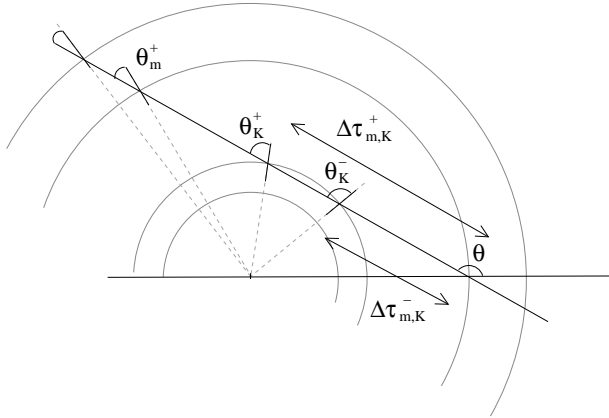


Fig. B.2. Definition of geometrical quantities.

Now, let us consider the case $\theta > \pi/2$, and designate K the innermost shell that is intersected by the ray. This shell is such that we have $\xi_{m,K} < \theta < \xi_{m,K-1}$, with the angle $\xi_{m,K}$ defined according to Eq. 6.

Manipulations of Eq. A.2 lead to

$$\Lambda^m = \sum_{k=K}^N \Lambda_{m,k} S^k \quad (\text{B.3})$$

with

– if $K+1 < k \leq m+1$:

$$\Lambda_{m,k} = \mathcal{A}_{m,k}(\theta) \quad (\text{B.4})$$

– if $m+1 < k \leq N$:

$$\Lambda_{m,k} = \sum_{l=-1}^1 c_{k+l}^{+l} e^{-\Delta\tau_{m,k+l}^+} \quad (\text{B.5})$$

– if $k = K+1$:

$$\Lambda_{m,K+1} = \mathcal{A}_{m,K+1}(\theta) + \alpha_K(\theta_K^+) e^{-\Delta\tau_{m,K}^+} + \gamma_K(\theta_K^-) e^{-\Delta\tau_{m,K}^-}$$

– if $k = K$:

$$\Lambda_{m,K} = \mathcal{A}_{m,K}(\theta) + \beta_K(\theta_K^+) e^{-\Delta\tau_{m,K}^+} + \left[(\alpha_K(\theta_K^-) + \beta_K(\theta_K^-)) e^{-\Delta\tau_{m,K}^-} + \gamma_K(\theta_K^+) e^{-\Delta\tau_{m,K}^+} \right] \times (1 - \delta(\theta - \xi_{m,K})) \quad (\text{B.6})$$

The quantity $\mathcal{A}_{m,k}(\theta)$ is defined according to

$$\mathcal{A}_{m,k}(\theta) = \sum_{l=-1}^1 \left[c_{k+l}^{+l}(\theta_{k+l}^+) e^{-\Delta\tau_{m,k+l}^+} + b_{k+l}^-(\theta_{k+l}^-) e^{-\Delta\tau_{m,k+l}^-} \right] \quad (\text{B.7})$$

$$\text{with } c_{k+l}^{+l} = \begin{cases} a_{k+l}^{+l} & \text{if } K < k+l < N \\ 0 & \text{otherwise} \end{cases}, \quad (\text{B.8})$$

$$\text{and } b_{k+l}^- = \begin{cases} a_{k+l}^- & \text{if } K < k+l \leq m \\ 0 & \text{otherwise} \end{cases}. \quad (\text{B.9})$$

In the expressions derived above, the dependence on the variable θ is not made explicit but would appear in the definitions of θ_{k+l}^\pm and $\Delta\tau_{m,k+l}^\pm$. The expressions obtained for $\Lambda_{m,k}$ are suitable for constructing an operator of whatever bandwidth. However, we are interested here in the diagonal of this operator:

– if $m > K+1$:

$$\Lambda_{m,m} = \sum_{l=-1}^1 c_{m+l}^{+l}(\theta_{m+l}^+) e^{-\Delta\tau_{m,m+l}^+} + \beta_m(\theta_m^-) + \gamma_{m-1}(\theta_{m-1}^-) e^{-\Delta\tau_{m,m-1}^-} \quad (\text{B.10})$$

– if $m = K+1$:

$$\Lambda_{m,m} = \sum_{l=0}^1 c_{m+l}^{+l}(\theta_{m+l}^+) e^{-\Delta\tau_{m,m+l}^+} + \beta_m(\theta_m^-) + \gamma_{m-1}(\theta_{m-1}^-) e^{-\Delta\tau_{m,m-1}^-} + \alpha_{m-1}(\theta_{m-1}^+) e^{-\Delta\tau_{m,m-1}^+} \quad (\text{B.11})$$

– if $m = K$:

$$\Lambda_{m,m} = \sum_{l=1}^1 c_{m+l}^{+l}(\theta_{m+l}^+) e^{-\Delta\tau_{m,m+l}^+} + \alpha_m(\theta_m^-) + \beta_m(\theta_m^-) + (\beta_m(\theta_m^+) + \gamma_m(\theta_m^+)) e^{-\Delta\tau_{m,m}^+} \quad (\text{B.12})$$

Hydrodynamics of giant planet formation

I. Overviewing the κ -mechanism

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Received September 15, 1996; accepted March 16, 1997

ABSTRACT

Context. To investigate the physical nature of the ‘nucleated instability’ of proto giant planets, the stability of layers in static, radiative gas spheres is analysed on the basis of Baker’s standard one-zone model.

Aims. It is shown that stability depends only upon the equations of state, the opacities and the local thermodynamic state in the layer. Stability and instability can therefore be expressed in the form of stability equations of state which are universal for a given composition.

Methods. The stability equations of state are calculated for solar composition and are displayed in the domain $-14 \leq \lg \rho / [\text{g cm}^{-3}] \leq 0$, $8.8 \leq \lg e / [\text{erg g}^{-1}] \leq 17.7$. These displays may be used to determine the one-zone stability of layers in stellar or planetary structure models by directly reading off the value of the stability equations for the thermodynamic state of these layers, specified by state quantities as density ρ , temperature T or specific internal energy e . Regions of instability in the (ρ, e) -plane are described and related to the underlying microphysical processes.

Results. Vibrational instability is found to be a common phenomenon at temperatures lower than the second He ionisation zone. The κ -mechanism is widespread under ‘cool’ conditions.

Key words. giant planet formation – κ -mechanism – stability of gas spheres

1. Introduction

In the *nucleated instability* (also called core instability) hypothesis of giant planet formation, a critical mass for static core envelope protoplanets has been found. Mizuno (1980) determined the critical mass of the core to be about $12 M_{\oplus}$ ($M_{\oplus} = 5.975 \times 10^{27} \text{ g}$ is the Earth mass), which is independent of the outer boundary conditions and therefore independent of the location in the solar nebula. This critical value for the core mass corresponds closely to the cores of today’s giant planets.

Although no hydrodynamical study has been available many workers conjectured that a collapse or rapid contraction will ensue after accumulating the critical mass. The main motivation for this article is to investigate the stability of the static envelope at the critical mass. With this aim the local, linear stability of static radiative gas spheres is investigated on the basis of Baker’s (1966) standard one-zone model.

Phenomena similar to the ones described above for giant planet formation have been found in hydrodynamical models concerning star formation where protostellar cores explode (Tscharnuter 1987, Balluch 1988), whereas earlier studies found quasi-steady collapse flows. The similarities in the (micro)physics, i.e., constitutive relations of protostellar cores and protogiant planets serve as a further motivation for this study.

2. Baker’s standard one-zone model

In this section the one-zone model of Baker (1966), originally used to study the Cepheid pulsation mechanism, will be briefly reviewed. The resulting stability criteria will be rewritten in terms of local state variables, local timescales and constitutive relations.

Baker (1966) investigates the stability of thin layers in self-gravitating, spherical gas clouds with the following properties:

- hydrostatic equilibrium,
- thermal equilibrium,
- energy transport by grey radiation diffusion.

For the one-zone-model Baker obtains necessary conditions for dynamical, secular and vibrational (or pulsational) stability (Eqs. (34a, b, c) in Baker 1966). Using Baker’s notation:

M_r	mass internal to the radius r
m	mass of the zone
r_0	unperturbed zone radius
ρ_0	unperturbed density in the zone
T_0	unperturbed temperature in the zone
L_{r0}	unperturbed luminosity
E_{th}	thermal energy of the zone

and with the definitions of the *local cooling time* (see Fig. 1)

$$\tau_{\text{co}} = \frac{E_{\text{th}}}{L_{r0}}, \quad (1)$$

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Fig. 1. Adiabatic exponent Γ_1 . Γ_1 is plotted as a function of \lg internal energy [erg g^{-1}] and \lg density [g cm^{-3}].

and the *local free-fall time*

$$\tau_{\text{ff}} = \sqrt{\frac{3\pi}{32G} \frac{4\pi r_0^3}{3M_r}}, \quad (2)$$

Baker's K and σ_0 have the following form:

$$\sigma_0 = \frac{\pi}{\sqrt{8}} \frac{1}{\tau_{\text{ff}}} \quad (3)$$

$$K = \frac{\sqrt{32}}{\pi} \frac{1}{\delta} \frac{\tau_{\text{ff}}}{\tau_{\text{co}}}; \quad (4)$$

where $E_{\text{th}} \approx m(P_0/\rho_0)$ has been used and

$$\delta = -\left(\frac{\partial \ln \rho}{\partial \ln T}\right)_P \quad (5)$$

$$e = mc^2$$

is a thermodynamical quantity which is of order 1 and equal to 1 for nonreacting mixtures of classical perfect gases. The physical meaning of σ_0 and K is clearly visible in the equations above. σ_0 represents a frequency of the order one per free-fall time. K is proportional to the ratio of the free-fall time and the cooling time. Substituting into Baker's criteria, using thermodynamic identities and definitions of thermodynamic quantities,

$$\Gamma_1 = \left(\frac{\partial \ln P}{\partial \ln \rho}\right)_S, \quad \chi_\rho = \left(\frac{\partial \ln P}{\partial \ln \rho}\right)_T, \quad \kappa_P = \left(\frac{\partial \ln \kappa}{\partial \ln P}\right)_T$$

$$\nabla_{\text{ad}} = \left(\frac{\partial \ln T}{\partial \ln P}\right)_S, \quad \chi_T = \left(\frac{\partial \ln P}{\partial \ln T}\right)_\rho, \quad \kappa_T = \left(\frac{\partial \ln \kappa}{\partial \ln T}\right)_T$$

one obtains, after some pages of algebra, the conditions for *stability* given below:

$$\frac{\pi^2}{8} \frac{1}{\tau_{\text{ff}}^2} (3\Gamma_1 - 4) > 0 \quad (6)$$

$$\frac{\pi^2}{\tau_{\text{co}} \tau_{\text{ff}}^2} \Gamma_1 \nabla_{\text{ad}} \left[\frac{1 - 3/4 \chi_\rho}{\chi_T} (\kappa_T - 4) + \kappa_P + 1 \right] > 0 \quad (7)$$

$$\frac{\pi^2}{4} \frac{3}{\tau_{\text{co}} \tau_{\text{ff}}^2} \Gamma_1^2 \nabla_{\text{ad}} \left[4\nabla_{\text{ad}} - (\nabla_{\text{ad}} \kappa_T + \kappa_P) - \frac{4}{3\Gamma_1} \right] > 0 \quad (8)$$

For a physical discussion of the stability criteria see Baker (1966) or Cox (1980).

We observe that these criteria for dynamical, secular and vibrational stability, respectively, can be factorized into

1. a factor containing local timescales only,
2. a factor containing only constitutive relations and their derivatives.

The first factors, depending on only timescales, are positive by definition. The signs of the left hand sides of the inequalities (6), (7) and (8) therefore depend exclusively on the second factors containing the constitutive relations. Since they depend only on state variables, the stability criteria themselves are *functions of the thermodynamic state in the local zone*. The one-zone stability can therefore be determined from a simple equation of state, given for example, as a function of density and temperature. Once the microphysics, i.e. the thermodynamics and opacities (see Table 1), are specified (in practice by specifying a chemical

Table 1. Opacity sources.

Source	$T/[\text{K}]$
Yorke 1979, Yorke 1980a	$\leq 1700^a$
Krügel 1971	$1700 \leq T \leq 5000$
Cox & Stewart 1969	$5000 \leq$

^a This is footnote a

composition) the one-zone stability can be inferred if the thermodynamic state is specified. The zone – or in other words the layer – will be stable or unstable in whatever object it is imbedded as long as it satisfies the one-zone-model assumptions. Only the specific growth rates (depending upon the time scales) will be different for layers in different objects.

We will now write down the sign (and therefore stability) determining parts of the left-hand sides of the inequalities (6), (7) and (8) and thereby obtain *stability equations of state*.

The sign determining part of inequality (6) is $3\Gamma_1 - 4$ and it reduces to the criterion for dynamical stability

$$\Gamma_1 > \frac{4}{3}. \quad (9)$$

Stability of the thermodynamical equilibrium demands

$$\chi_\rho > 0, \quad c_v > 0, \quad (10)$$

and

$$\chi_T > 0 \quad (11)$$

holds for a wide range of physical situations. With

$$\Gamma_3 - 1 = \frac{P}{\rho T} \frac{\chi_T}{c_v} > 0 \quad (12)$$

$$\Gamma_1 = \chi_\rho + \chi_T (\Gamma_3 - 1) > 0 \quad (13)$$

$$\nabla_{\text{ad}} = \frac{\Gamma_3 - 1}{\Gamma_1} > 0 \quad (14)$$

we find the sign determining terms in inequalities (7) and (8) respectively and obtain the following form of the criteria for dynamical, secular and vibrational *stability*, respectively:

$$3\Gamma_1 - 4 =: S_{\text{dyn}} > 0 \quad (15)$$

$$\frac{1 - 3/4 \chi_\rho}{\chi_T} (\kappa_T - 4) + \kappa_P + 1 =: S_{\text{sec}} > 0 \quad (16)$$

$$4\nabla_{\text{ad}} - (\nabla_{\text{ad}} \kappa_T + \kappa_P) - \frac{4}{3\Gamma_1} =: S_{\text{vib}} > 0. \quad (17)$$

The constitutive relations are to be evaluated for the unperturbed thermodynamic state (say (ρ_0, T_0)) of the zone. We see that the one-zone stability of the layer depends only on the constitutive relations Γ_1 , ∇_{ad} , χ_T , χ_ρ , κ_P , κ_T . These depend only on the unperturbed thermodynamical state of the layer. Therefore the above relations define the one-zone-stability equations of state S_{dyn} , S_{sec} and S_{vib} . See Fig. 2 for a picture of S_{vib} . Regions of secular instability are listed in Table 1.

Fig. 2. Vibrational stability equation of state $S_{\text{vib}}(\lg e, \lg \rho) > 0$ means vibrational stability.

3. Conclusions

1. The conditions for the stability of static, radiative layers in gas spheres, as described by Baker's (1966) standard one-zone model, can be expressed as stability equations of state. These stability equations of state depend only on the local thermodynamic state of the layer.
2. If the constitutive relations – equations of state and Rosseland mean opacities – are specified, the stability equations of state can be evaluated without specifying properties of the layer.
3. For solar composition gas the κ -mechanism is working in the regions of the ice and dust features in the opacities, the H_2 dissociation and the combined H, first He ionization zone, as indicated by vibrational instability. These regions of instability are much larger in extent and degree of instability than the second He ionization zone that drives the Cepheid pulsations.

Acknowledgements. Part of this work was supported by the German *Deutsche Forschungsgemeinschaft*, DFG project number Ts 17/2–1.

References

- Baker, N. 1966, in *Stellar Evolution*, ed. R. F. Stein, & A. G. W. Cameron (Plenum, New York) 333
- Balluch, M. 1988, *A&A*, 200, 58
- Cox, J. P. 1980, *Theory of Stellar Pulsation* (Princeton University Press, Princeton) 165
- Cox, A. N., & Stewart, J. N. 1969, *Academia Nauk, Scientific Information* 15, 1
- Mizuno H. 1980, *Prog. Theor. Phys.*, 64, 544
- Tscharnutter W. M. 1987, *A&A*, 188, 55
- Terlevich, R. 1992, in *ASP Conf. Ser. 31, Relationships between Active Galactic Nuclei and Starburst Galaxies*, ed. A. V. Filippenko, 13
- Yorke, H. W. 1980a, *A&A*, 86, 286
- Zheng, W., Davidsen, A. F., Tytler, D. & Kriss, G. A. 1997, preprint